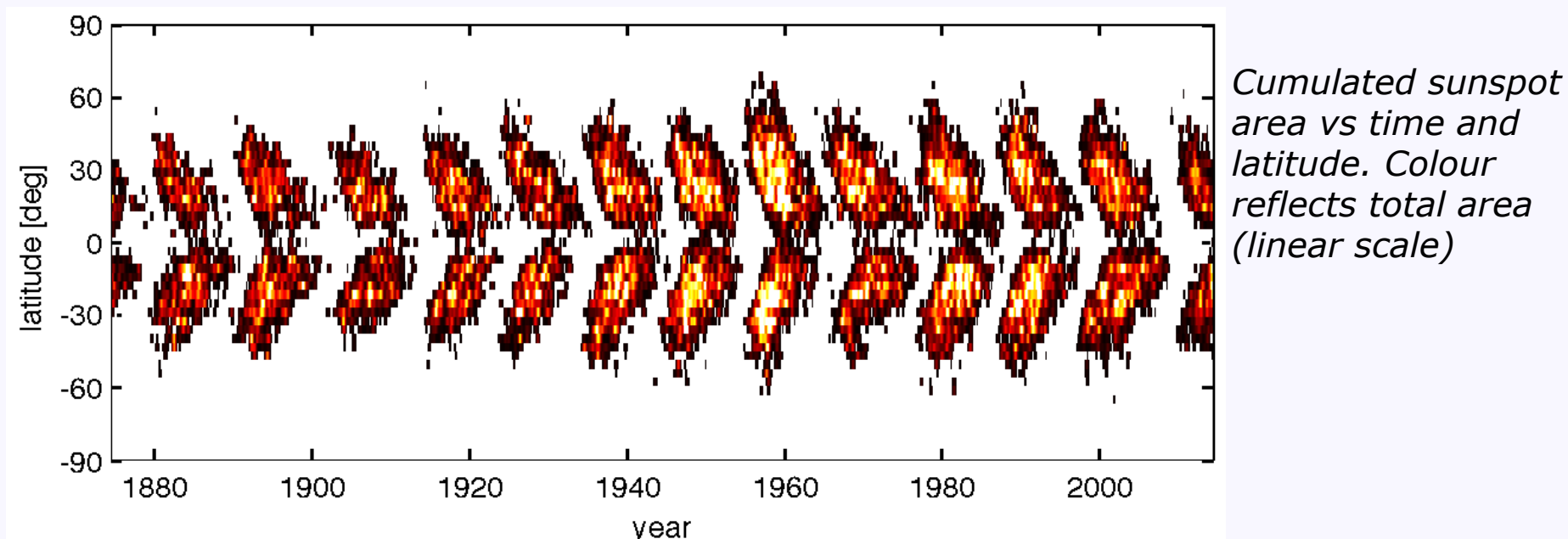


# The butterfly diagram: from a phase space portrait to a predator-prey model

Thierry Dudok de Wit  
LPC2E, CNRS and University of Orléans

## 1. The butterfly diagram: still a lot to discover

The latitudinal distribution of the area of sunspots (aka, the *butterfly diagram*) is much more informative than the sunspot number as it reflects the way sunspots migrate under the effect of the solar dynamo.



There have been **many attempts to reduce this spatio-temporal diagram** to sets of simpler proxies of the solar dynamo: projection of spherical coordinates, principal component analysis, etc [Gokhale, Knaack, Hathaway, Mininni, Consolini, ...] but their physical interpretation is often debatable.

**Our objective:** use blind source separation to reduce the butterfly diagram to proxies that have a more immediate physical meaning.

## 2. Data reduction

Hypothesis: decompose the sunspot areas into separable « modes »

$$A(t, \theta) = \sum_k S_k(\theta) \cdot M_k(t) + \epsilon(t, \theta)$$

residual error

sunspot area

latitudinal profile, or « source »

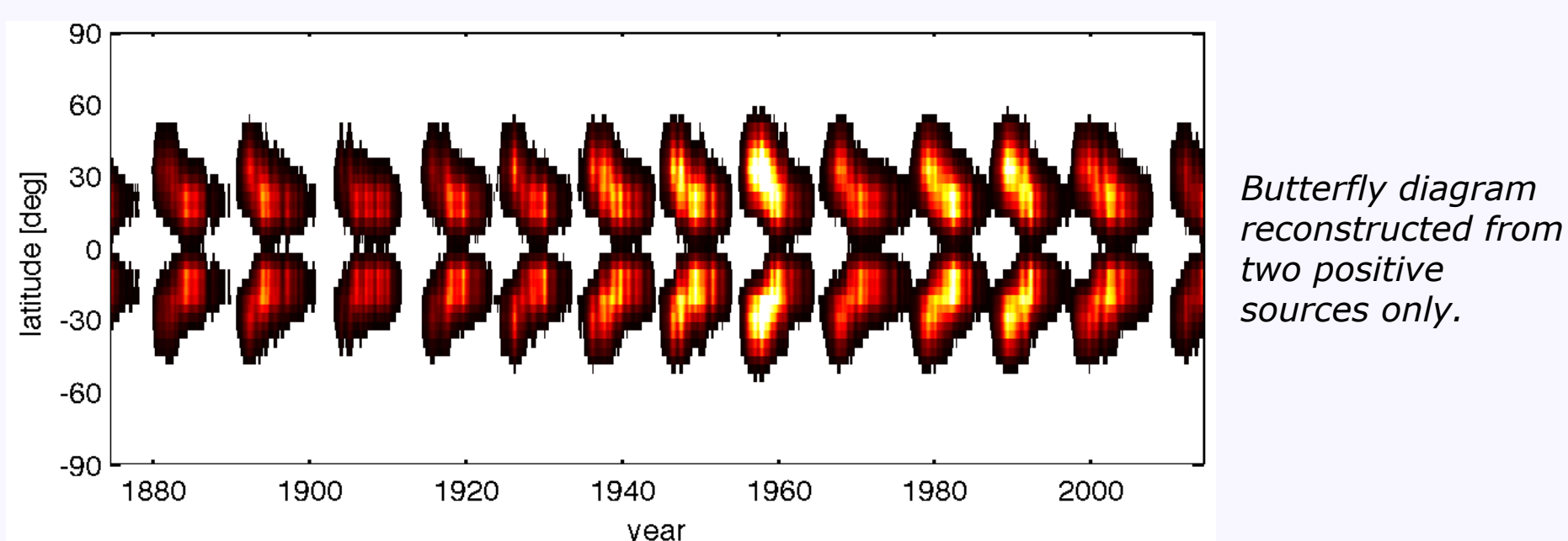
temporal profile, or « mixing coefficient »

Principal component analysis tells us that all the salient features of the butterfly diagram can be captured by just 2 to 3 modes. But principal components can be negative = no physical meaning.

We apply instead a Bayesian Positive Source Separation [Moussaoui et al., 2002] technique, and constrain the modes with:

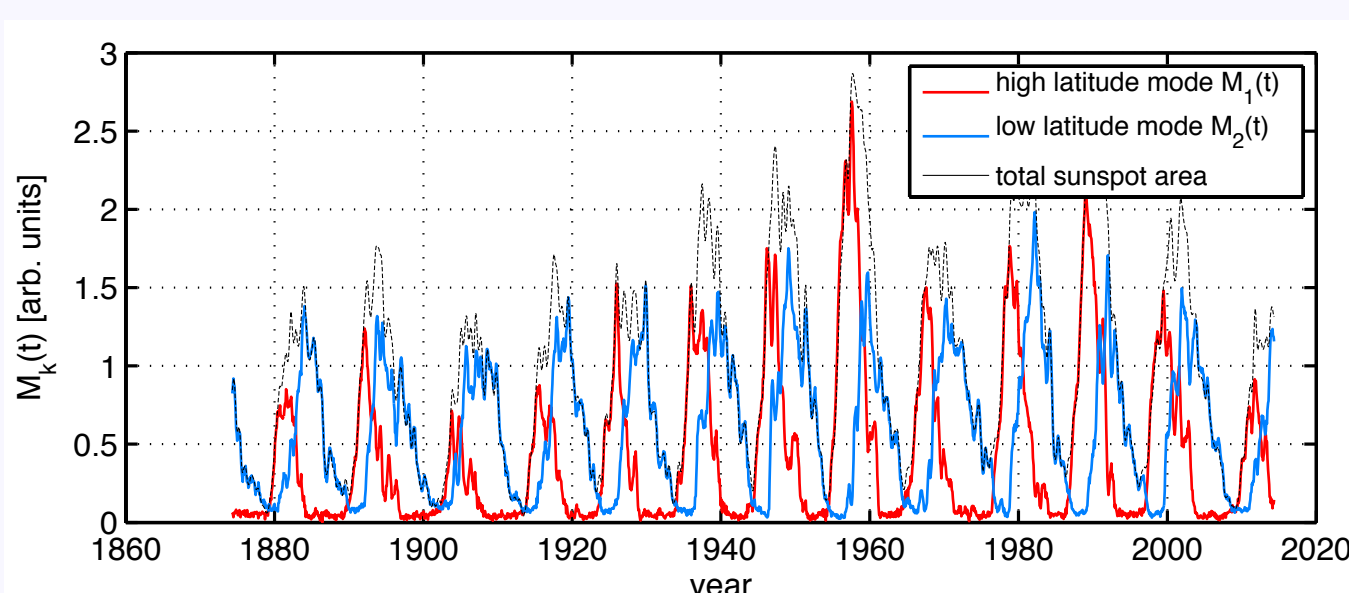
- temporal profiles  $M_k(t)$  must be independent
- temporal profiles  $M_k(t)$  and sources  $S_k(\theta)$  must be  $\geq 0$

**2 sources only suffice to capture all the coherent features of the butterfly diagram.** Additional sources merely describe small-scale fluctuations.

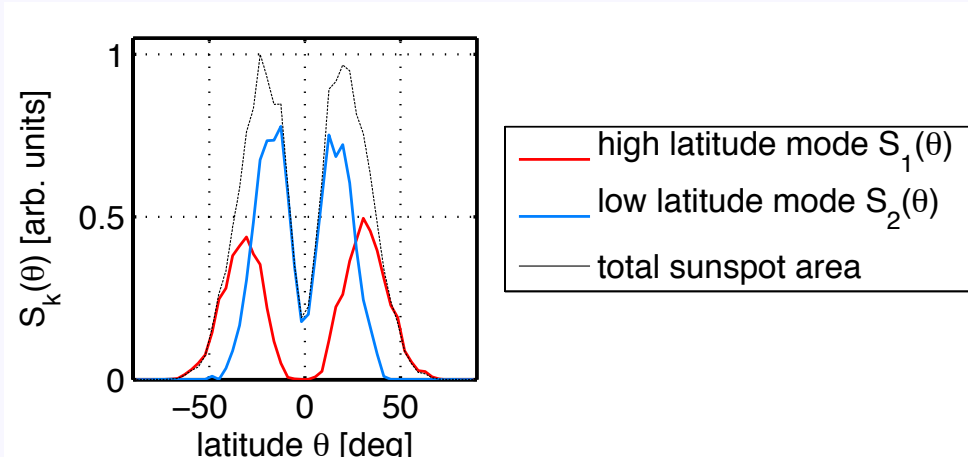


The key properties of the solar cycle (migration speed, amplitude, ...) are now captured by our two temporal profiles  $M_1(t)$  and  $M_2(t)$ .

## 3. What the modes look like



The two temporal profiles  $M_1$  and  $M_2$ .



The corresponding sources, or latitudinal profiles  $S_1$  and  $S_2$ .

## 4. Interpretation

The **high latitude mode  $M_1(t)$**  describes the emergence of sunspots at high latitudes and is representative of the conversion from poloidal to toroidal flux. This mode directly feeds  $M_2$  = it is the « prey »

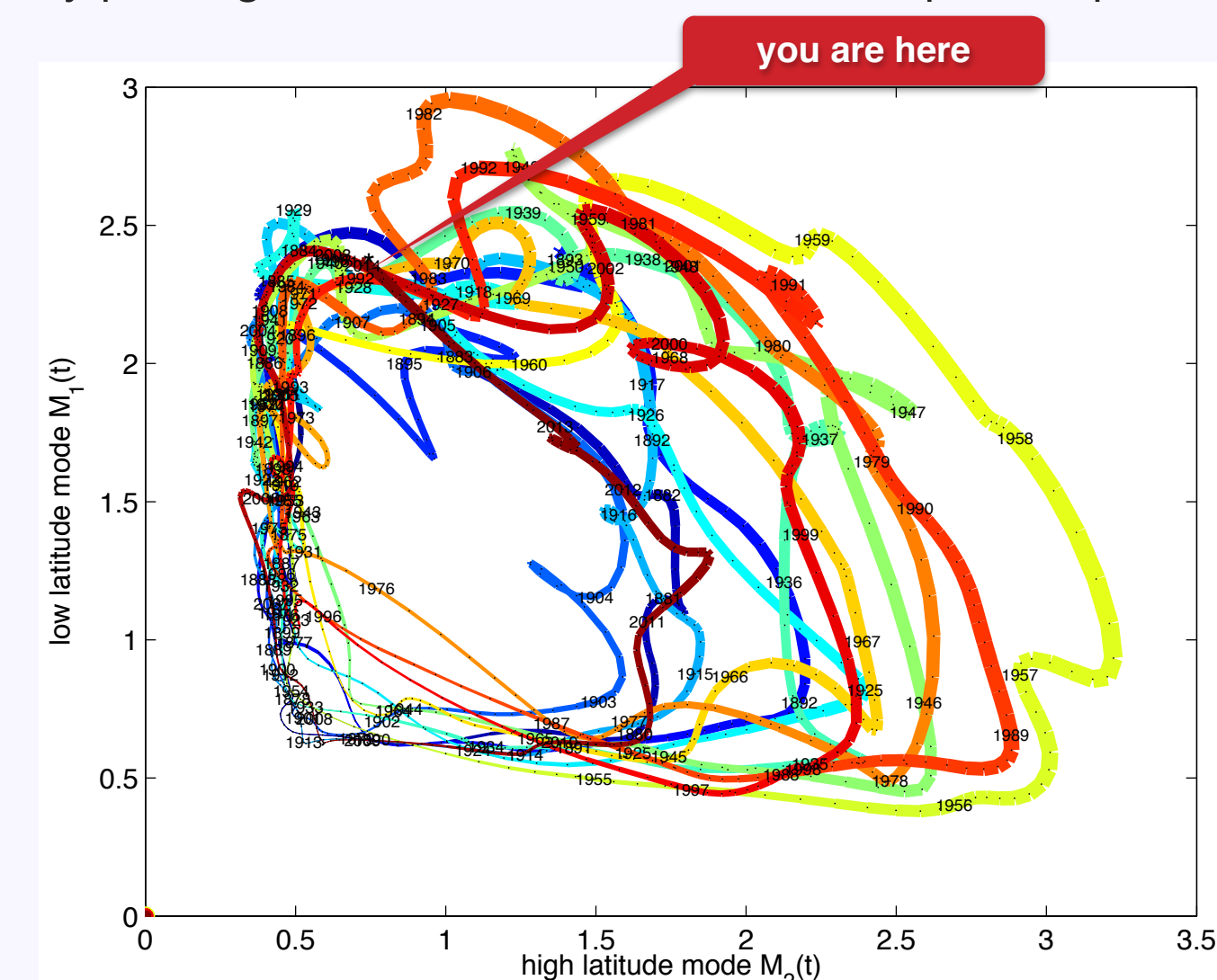
The **low latitude mode  $M_2(t)$**  describes the disappearance of sunspots at low latitudes and is representative of the conversion from toroidal to poloidal flux. This mode is the « predator » of  $M_1$ , and subsequently feeds the next cycle.

This representation (re)opens several perspectives

- Describe the asymmetry between both hemispheres and relate this to their synchronisation.
- Asses the evidence for deterministic vs stochastic behavior
- Understand how the characteristics of each cycle (drift speed toward equator, amplitude, duration, etc.) are related.
- Occurrence of Gnevyshev gap at transition from one mode to the other. etc

## 5. Phase space representation

By plotting  $M_2$  vs  $M_1$  we obtain a concise phase space representation.



Phase space plot. Colour reflects time and line width the total sunspot area. The data have been smoothed over 4 months to ease visualisation.

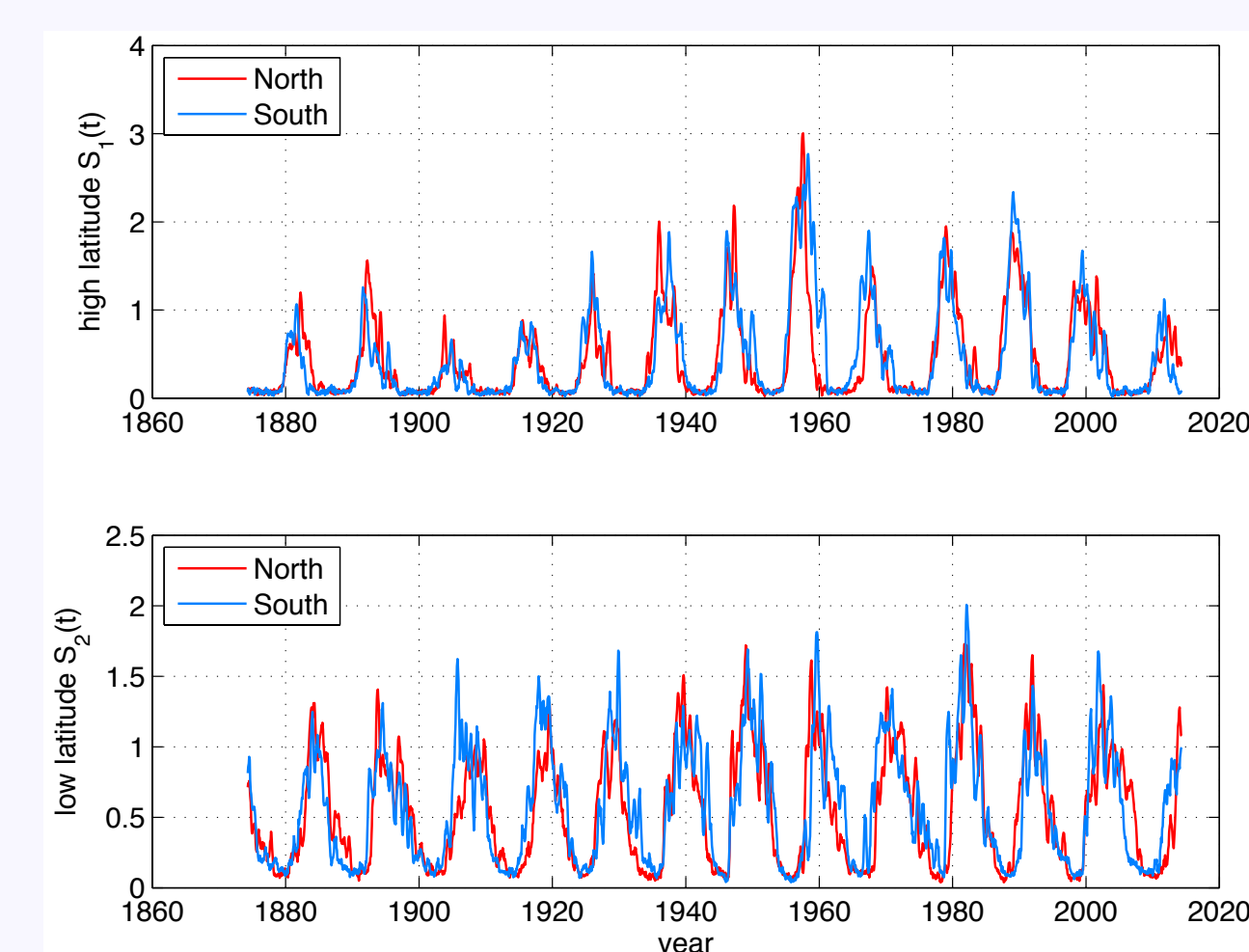
We take beforehand the square root of the sunspot area to stabilise its variance (Anscombe transform)

### Interpretation

- these orbits are reminiscent of the Lotka-Volterra predator-prey model, which thus gives us a simple analogy of the butterfly diagram.
- two solar cycles are similar **only** if their orbits overlap: we find that the last cycle (nr 23) is analogous to the one that peaked in 1883, and not to the one of 1914, as often suggested.
- this plot gives deep insight in how the transition during sunspot minimum affects the subsequent cycle. More on this soon!

## 6. Hemispheric asymmetries ?

By estimating the modes separately from both hemispheres, we get a detailed picture of how these asymmetries actually are.



High latitude mode  $M_1$  estimated separately for both hemispheres

Low latitude mode  $M_2$  estimated separately for both hemispheres

Acknowledgements : This study received funding from the European Community's Seventh Framework Programme (FP7-SPACE-2012-2) under the grant agreement nr. 313188 (SOLID) and COST action TOSCA. I gratefully thank the RGO for making their data available.

